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On the morphology of the compact dust shell in the symbiotic system HM Sagittae^{★,★★} (Research Note)

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ABSTRACT

Context. The symbiotic system HM Sagittae consists of a Mira star and a secondary White Dwarf component. The dust content of the system was severely affected by the nova outburst in 1975, which is still ongoing. The capabilities of optical interferometry operating in the mid-IR allow us to investigate the current geometry of the dust envelope.

Aims. We test our previous spectro-interferometric study of this system with new interferometric configurations, increasing the uv coverage and allowing us to ascertain the appearance of the source between 8 and 13 μm .

Methods. We used the MIDI instrument of the VLTI with the unit telescopes (UTs) and auxiliary telescopes (ATs) providing baselines oriented from $\text{PA}=42^\circ$ to 127° . The data are interpreted by means of an elliptical Gaussian model and the spherical radiative transfer code DUSTY.

Results. We demonstrate that the data can be reproduced well by an optically thick dust shell of amorphous silicate, typical of those encountered around Mira stars, whose measured dimension increases from 8 to 13 μm . We confirm that the envelope is more extended in a direction perpendicular to the binary axis. The level of elongation increases with wavelength in contrast to our claim in a previous study.

Conclusions. The wider uv coverage allows us to deepen our previous investigations of the close circumstellar structure of this object.

Key words. techniques: interferometric - techniques: high angular resolution - stars: AGB and post-AGB - stars: binaries - stars: symbiotic - stars: circumstellar matter - stars: mass-loss

1. Introduction

HM Sge is a D-type (dust dominated IR) symbiotic system that erupted as a symbiotic nova in 1975 (Dokuchaeva, 1976), evolving from a 17th to a 11th magnitude star with a rich emission-line spectrum. The cool component is a Mira star and the hot component is a White Dwarf (WD), which had escaped detection being enshrouded in the dense Mira envelope before this dramatic event.

Previous works by Schild et al. (2001) and Bogdanov & Taranova (2001) were published on HM Sge the same year, evaluating the parameters of the dust in the system by fitting the ISO/SWS data using the DUSTY radiative transfer code (Ivezić

et al., 1999). Both groups used single and double component models but did not reach identical conclusions. Schild et al. (2001) favored a 2-shell model, which accounted for the effect of the White Dwarf on the dust, whereas Bogdanov & Taranova (2001) proposed a single-shell model. Since the extensions of the dusty shell differ from one model to another, mid-infrared interferometry data enabled Sacuto et al. (2007) to determine the most suitable of the proposed models. The authors showed that the morphology of the dust was accounted for well by the single shell model of Schild et al. (2001), which included a compact and optically thick circumstellar environment, typical of a Mira star. However, because of its high optical depth, the model is inconsistent with the presence of non-absorbed emission of the Mira in the near-IR, and Sacuto et al. (2007) indicated that this issue could be solved by assuming a more complex spatial distribution of dust. We attempt to confirm the morphology of the dusty envelope using observations acquired with five new interferometric baselines secured in 2007.

* Based on observations made with the Very Large Telescope Interferometer at Paranal Observatory under programs 075.D-0484, 077.D-0216, and 079.D-0213

** Reduced visibilities and differential phases are available in electronic form at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsweb.u-strasbg.fr/cgi-bin/qcat?J/A+A/?/?>

2. Observations

The Very Large Telescope Interferometer (VLTI) of ESO's Paranal Observatory was used with MIDI, the MID-infrared Interferometric recombiner (Leinert et al., 2003). MIDI combines the light of two telescopes and provides spectrally resolved visibilities in the N band atmospheric window.

The first observations of HM Sge were conducted in 2005 and 2006 with the VLT unit telescopes (UTs) UT2, UT3, and UT4, providing projected baselines and projected angles, in the range of 32-59 meters and 42° to 105° respectively (see the first two parts of Table 1). An interpretation of these data is described in Sacuto et al. (2007).

The acquisition images enabled us to confirm that the mid-IR source was unresolved by the data from the individual UT. This implied that most of the mid-infrared flux originated in the inner 300 mas of the source, corresponding to the Airy pattern of the UTs and defining the Field Of View (FOV) of the interferometric observations. This also means that the flux detected within the $1.1''$ FOV of the VLTI auxiliary telescopes (ATs), originates in this inner region. Therefore, the visibilities extracted from the ATs and the UTs are fully consistent with each other.

Complementary observations were completed in 2007 with the ATs A0, D0, G1, and H0, providing projected baselines and projected angles, in the range of 64-89 meters and 72° to 127° respectively (see the third part of Table 1). With longer projected baselines, these configurations allow us to probe smaller circumstellar regions. The new projected angles provide additional constraints about the morphology of the dust shell.

Table 1 presents the journal of interferometric observations. The calibrators, HD188512 (G8IV, $\text{diam}=1.98\pm0.02$ mas), HD187642 (A7V, $\text{diam}=3.22\pm0.01$ mas), HD187076 (M2II, $\text{diam}=8.05\pm0.13$ mas), HD167618 (M3.5III, $\text{diam}=11.33\pm0.04$ mas), and HD206778 (K2Ib, $\text{diam}=8.38\pm0.09$ mas)¹, were observed immediately before or after each science target observation. The third part of the table gathers the additional observations of the object obtained in April, May, and July 2007.

Chopped acquisition images were recorded ($f=2\text{Hz}$, 2000 frames, 4 ms per frame) to ensure the accurate acquisition of the target. The acquisition filter was an N-band filter. Photometry was obtained before and after the interferometric observations with the HIGH-SENS mode of MIDI, using the prism that provides a spectral dispersion of about 30. The data reduction software packages² MIA and EWS (Jaffe, 2004) were used to prepare the spectra and visibilities (Chesneau et al., 2005; Ratzka et al., 2007). MIA is based on a power spectrum analysis and uses a fast Fourier transformation (FFT) to calculate the Fourier amplitude of the fringe packets,

Table 1. Journal of all available MIDI observations. The calibrators used to calibrate the visibilities are given below the science target. The phase of the Mira (φ_{Mira}) during the observations is indicated. The configuration used for the observations is given. The length and position angle of the projected baseline are also indicated.

Star	UT date & Time	φ_{Mira}	Config.	Base[m]	PA[deg]
HM Sge	2005-07-23 06:48:42	0.75	U2-U3	46.5	44
HD188512	2005-07-23 07:12:10	...	-
HM Sge	2005-07-24 02:41:11	0.75	U2-U3	32.1	42
HD188512	2005-07-24 02:58:55	...	-
HM Sge	2005-07-24 03:42:18	0.75	U2-U3	37.8	47
HD188512	2005-07-24 04:03:14	...	-
HM Sge	2005-07-24 06:14:49	0.75	U2-U3	46.2	47
HD188512	2005-07-24 06:37:58	...	-
HM Sge	2006-05-17 09:04:02	0.31	U3-U4	59.3	105
HD187642	2006-05-17 09:27:14	...	-
HM Sge	2006-06-11 09:05:09	0.36	U3-U4	46.8	101
HD187642	2006-06-11 08:42:52	...	-
HM Sge	2007-04-04 09:39:22	0.93	A0-G1	89	118
HD187076	2007-04-04 09:06:17	...	-
HM Sge	2007-05-08 08:52:55	0.99	G1-D0	66	127
HD167618	2007-05-08 10:16:28	...	-
HM Sge	2007-05-10 08:34:58	0.99	H0-D0	60	78
HD187076	2007-05-10 08:55:17	...	-
HM Sge	2007-05-16 09:48:33	0.00	H0-D0	64	72
HD206778	2007-05-16 09:20:11	...	-
HM Sge	2007-07-30 05:54:54	0.15	A0-G1	65	107
HD206778	2007-07-30 05:23:12	...	-

while EWS uses a shift-and-add algorithm in the complex plane, averaging appropriately modified individual exposures (dispersed channeled spectra) to obtain the complex visibility. An extensive study of the fluctuations of the instrumental transfer function was carried out in all periods. The rms scatter of these calibration measurements was ~ 0.05 between $8\text{ }\mu\text{m}$ and $13\text{ }\mu\text{m}$, corresponding to a relative error of 10%.

All visibility and differential phase data as well as all the characteristics of the observations are available from the CDS (Centre de Données astronomiques de Strasbourg); all data products are stored in the FITS-based, optical-interferometry data-exchange-format (OI-FITS) described by Pauls et al. (2005).

In Fig. 1, the MIDI spectra for each period (2005-2006-2007) is compared with the ISO/SWS spectra (Sloan et al., 2003) acquired at two different phases of the Mira. The level of flux received by MIDI is equivalent to that of ISO, implying that most of the N flux is concentrated within the 300 mas beam of a 8 m telescope. It also means that the level of emission has probably not changed significantly since the ISO observations (1996-1997), keeping in mind a typical error for the absolute SWS flux calibration of 15 to 30% (Schaeidt et al., 1996). For physical and technical reasons, the flux error bars of the photometry from the UTs are below the 10% level, whereas the accuracy of the photometry extracted from the ATs is much

¹ MIDI calibrator database

² <http://www.mpia-hd.mpg.de/MIDISOFT/>,
<http://www.strw.leidenuniv.nl/~nevec/MIDI/>

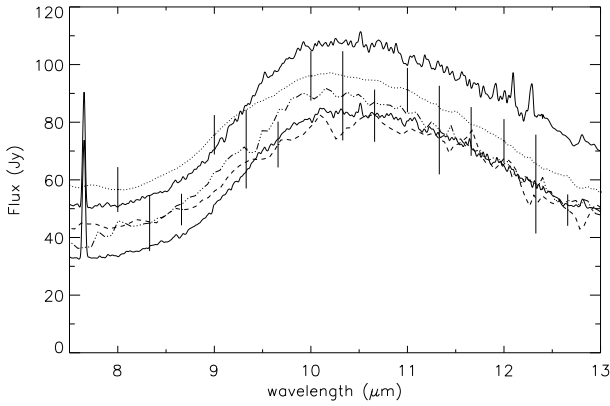


Fig. 1. MIDI and ISO/SWS flux of the source in the mid-infrared. Dotted line with $\sim 10\%$ error bars is the 2005 MIDI flux from the UTs ($\phi \sim 0.75$); dashed line with $\sim 10\%$ error bars is the 2006 MIDI flux from the UTs ($\phi \sim 0.30$); dashed dot dot line with $\sim 25\%$ error bars is the 2007 MIDI flux from the ATs ($\phi \sim 0.00$). The two solid lines are the ISO/SWS spectra (upper one: 1996, $\phi = 0.65$; bottom one: 1997, $\phi = 0.08$)

worse with an error bar of 25% (see Fig. 1). The error levels are estimated from the standard deviation of the flux-calibrated spectra determined from each night and each telescope, *i.e.* typically 4 to 8 spectra. As a consequence, it is difficult to differentiate between photometric variations in one phase from another. The number of measurements acquired in each period is also limited. Therefore, we decided in the following to continue the spectro-interferometric study of the star neglecting its variability.

3. Radiative transfer modeling

The spectro-interferometric observations are interpreted in terms of spherical models generated by the DUSTY package (Ivezić et al., 1999). We first fit the spectrophotometric data of the star, then we determine the corresponding synthetic spectrally-dispersed visibility profiles throughout the N band (8-13 μm) and compare them with the MIDI visibilities for each baseline.

3.1. The compact single-dust-shell model

The best-fit models obtained in our previous work (see Sect. 5 of Sacuto et al. (2007)) involved only a single shell component around the Mira. The modeling was based on the simultaneous fit of the ISO/SWS spectrometric (Sloan et al., 2003) and interferometric data for periods 2005 and 2006 (see the first two part of Table 1). Different models were obtained: one involving pure amorphous silicates, the other including carbon material in the dusty layers. We tested these models upon the new MIDI interferometric measurements (see the third part of Table 1) and also some near-infrared photometric data not used in the first study.

Near-infrared photometric data provide an additional constraint

Table 2. Parameters for HMsge deduced from the best-fit DUSTY model to the spectrophotometric and interferometric data.

Parameter	Value
Effective temperature (K)	3000
Central star diameter (mas)	1.6
Luminosity (L_{\odot})	5000
Distance (kpc)	1.5
Shell inner radius: ϵ_{in} (mas)	3.6
Inner boundary temperature (K)	1600
Grain chemical composition	100% Sil-Ow ¹
Density power law coefficient	2.2
Grain size distribution	MRN ²
Geometrical thickness (ϵ_{in})	1000
Visual optical depth	25.2
10 μm optical depth	2.15

¹Sil-Ow stands for 'warm' silicates of Ossenkopf et al. (1992).

²grain size distribution as described by Mathis et al. (1977).

of the radiation originating in the Mira star, whereas the interferometric data provide the necessary high angular resolution to probe the internal layers and to derive information about the morphology of the structure.

A least square fitting minimization was performed taking into account spectrophotometric and interferometric data. This minimization was completed for a large grid of parameter values by varying the effective temperature of the star from 2800 to 3200 K ($\Delta T_{\text{eff}} = 100$ K), the inner boundary temperature from 1400 to 1800 K ($\Delta T_{\text{in}} = 100$ K), the density power-law coefficient from 1.8 to 2.4 ($\Delta p = 0.1$), and the 10 μm optical depth from 1.5 to 2.5 ($\Delta \tau_{10\mu\text{m}} = 0.05$).

Figure 2 shows the best-fit model allowing a simultaneous fit of spectrophotometric and interferometric data of the star. Table 2 presents the parameters of HMsge deduced from this fit.

The results do not differ significantly from the parameters of the best-fit model deduced in the previous work of Sacuto et al. (2007), and reproduce well the near-infrared photometric data and the new interferometric data of the star. However, the new data enable us to exclude the model including carbon dust proposed to improve the fit. This assumption, involving a large fraction of carbon, is not physically justified given the spectral type of the dust-forming source, an M-type Mira star. In the absence of any clear signature of other dust species and given the spatial complexity of the source, it is extremely difficult to use the dust composition as a parameter for improving the fit.

The interpretation by means of a single dust shell surrounding the Mira star was first rejected by Schild et al. (2001) because of its high opacity, which was inconsistent with the detection of the Mira photometric cycle in the near-infrared. However, interferometric data imply that a simple dust shell is the most suitable solution.

With the help of these new observations, the increase of the uv coverage allows us now to investigate the evidence of departure of the circumstellar environment from spherical symmetry.

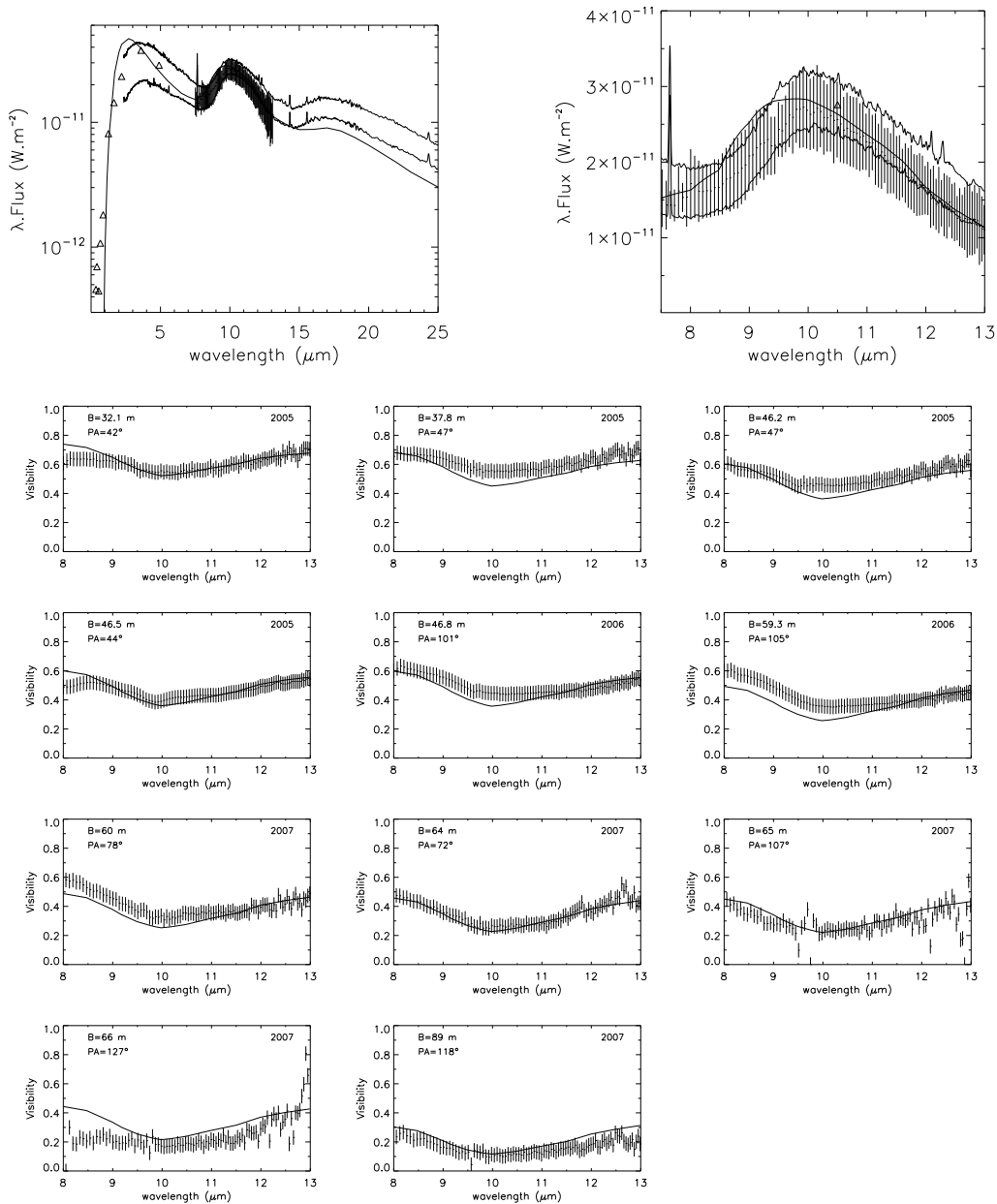


Fig. 2. *Top-left:* Best-fit model of the single dust shell spectrum (thin line) superimposed on the ATs MIDI flux (error bars), the 2 ISO/SWS spectra (upper thick line: 1996, $\phi=0.65$, and bottom thick line: 1997, $\phi=0.08$), and the photometric data (open triangles) with (U/B/V/R/J/H/K/L/M/N) data from Yudin et al. (1994) and *I* band from Chochol et al. (2004). *Top-right:* close-up view of the best-fit model to the dust feature. *Bottom:* the corresponding model visibility (solid line) superimposed on the MIDI visibilities (error bars) for the eleven projected baselines.

4. Morphological interpretation

The phase shifts or differential phases is information extracted from the difference between the phase at a given wavelength and the mean phase determined in the full N-band region. It provides partial information (depending on the uv coverage) about the asymmetry of the source, in addition to the information contained in the visibility amplitudes. As an example, strong phase signals were detected in sources harboring dusty disks (Deroo et al., 2007; Ohnaka et al., 2008). Previous results of differential phases obtained by Sacuto et

al. (2007) indicated no obvious signature of any asymmetry in the close environment of the star. Due to the low rms ($< \pm 5^\circ$) of the calibrated differential phases averaged over all the projected baselines, new interferometric configurations confirm the global centrosymmetry of the object in the vicinity of the source (see Fig. 3). This allows us to justify in the following the utilization of a centrosymmetric geometrical model.

The geometrical constraints obtained from the data presented in Sacuto et al. (2007) show that the circumstellar environment of the star is elongated perpendicularly to the binary

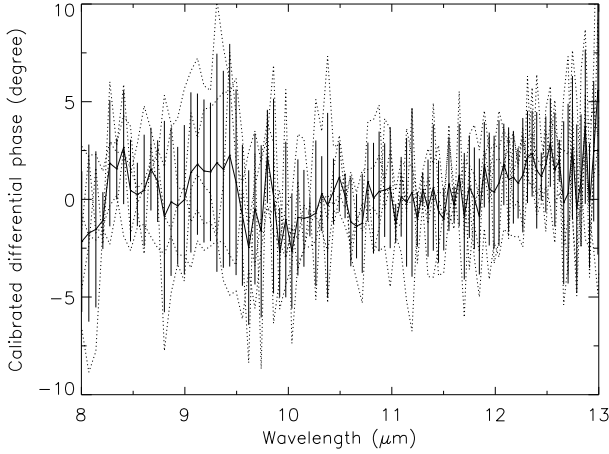


Fig. 3. Calibrated differential phase of the object evaluated from all the determined ATs projected baselines (dotted lines). Error bars correspond to the rms of the differential phase averaged over all the projected baselines. The bold solid line is the central value of these error bars.

axis (PA=130° NE; Eyres et al. 2001), in the direction of the Raman line polarization (Schmid et al., 2000). This information was obtained by comparing the half width at half maximum of the equivalent one-dimensional Gaussian distribution from various projected baseline position angles. With the new baselines, we can perform a more robust analysis using a two-dimensional fitting in the Fourier space by means of an elliptical Gaussian flux distribution, appropriate for studying deviations from spherical symmetry.

4.1. Elliptical Gaussian model

The use of purely geometrical models can help to constrain the object morphology. Because the circumstellar envelope is optically thick ($\tau_{0.55\mu\text{m}} = 25.2$; see Table 2), a Gaussian brightness distribution centered on the star is a natural approximation of the flux distribution in the plane of the sky.

The visibility of the elliptical Gaussian model can be expressed by (Tycner et al., 2004)

$$V(u', v') = \exp \left[\frac{-\pi^2 \theta_{\text{mj}}^2 (r^2 u'^2 + v'^2)}{4 \ln 2} \right] \quad (1)$$

where θ_{mj} is the angular size FWHM of the major axis, r is the axial ratio of the minor to the major axis, and (u', v') are the modified spatial frequencies. These frequencies can be expressed in terms of the spatial frequencies (u, v) by using a co-ordinate transformation

$$u' = u \cos \phi - v \sin \phi \quad (2)$$

$$v' = u \sin \phi + v \cos \phi \quad (3)$$

where $u = \frac{B_p}{\lambda} \cos \theta_p$ and $v = \frac{B_p}{\lambda} \sin \theta_p$, where B_p and θ_p are the projected baseline and the projected angle on the sky

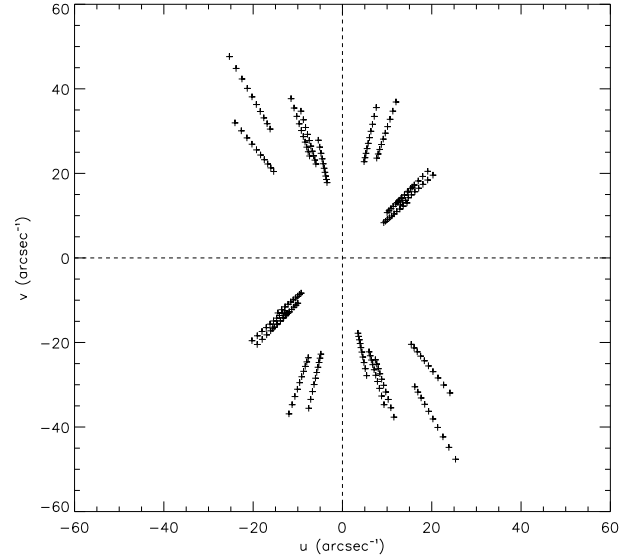


Fig. 4. Spectrally-dispersed uv coverage of the 2005-2006-2007 observations from 8 to 12.5 μm .

respectively, and λ is the observed wavelength. Finally, ϕ is the position angle (measured east from north) of the major axis.

By completing a non-linear least squares fitting, based on the Levenberg-Marquardt method, we obtained the best-fit elliptical Gaussian model to the data with a reduced χ^2 of 0.27. Figure 4 shows the spectrally-dispersed uv coverage of all the observations (see Table 1). Figure 5 provides the brightness distribution of the best-fit elliptical Gaussian model of HM Sge between 8 and 12.5 μm . Figure 6 shows the resulting FWHM of the major axis, and the ratio of the minor to the major axis for the best-fit elliptical Gaussian model as a function of wavelength.

The best-fit model is obtained for an extension of the structure perpendicular to the binary axis located at PA=130° NE (Eyres et al., 2001) as presented in our previous work. The length of the major axis is also approximatively the same as that published in Sacuto et al. (2007). However, the level of elongation *increases* from 8 to 12.5 μm , in contrast to results found before. This difference can be explained by the wider uv coverage allowing the geometry of the structure to be more tightly constrained.

5. Conclusions

The wider uv coverage data has allowed us to confirm that the dust envelope around the Mira is described well by an optically thick dust shell that consists predominantly of amorphous silicates. We also confirmed the longer extension of the circumstellar envelope in a direction perpendicular to the binary axis, and the increase in the size of the structure between 8 and 12.5 μm . It has also been demonstrated that the amount of elongation is higher at 12.5 μm and steadily decreases toward 8 μm , but the ratio between the two axis remains limited at most to 0.85, implying only a slight departure from spherical symmetry. This behavior can naturally be interpreted. The layers emit-

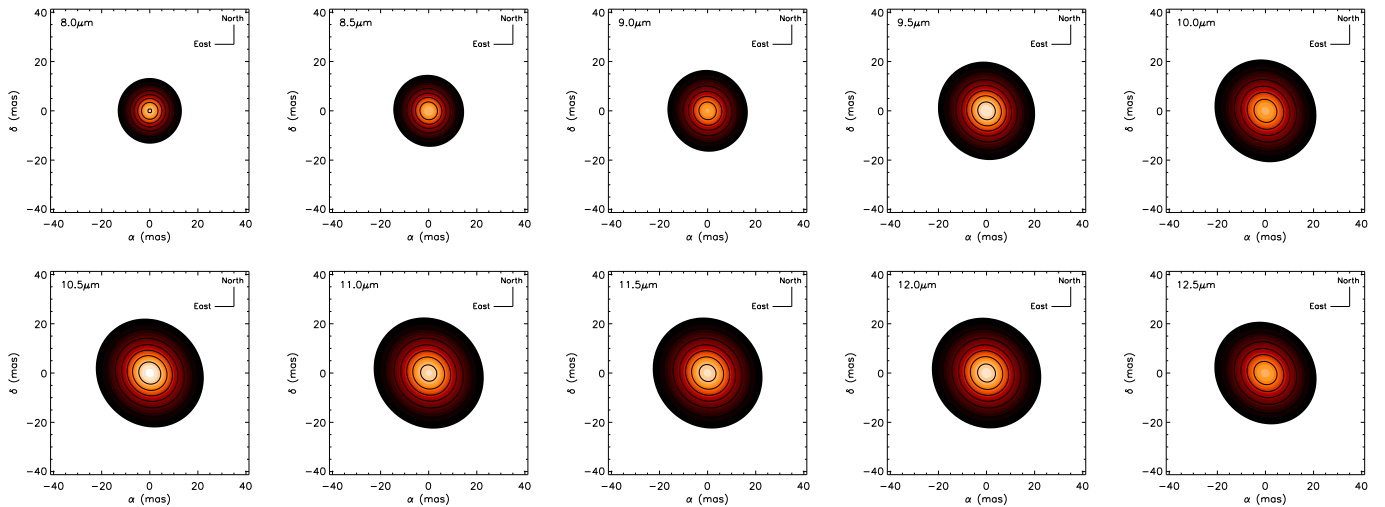


Fig. 5. Brightness distribution of the best-fit elliptical Gaussian model of HM Sge from 8 to 12.5 μm (north is up and east is left).

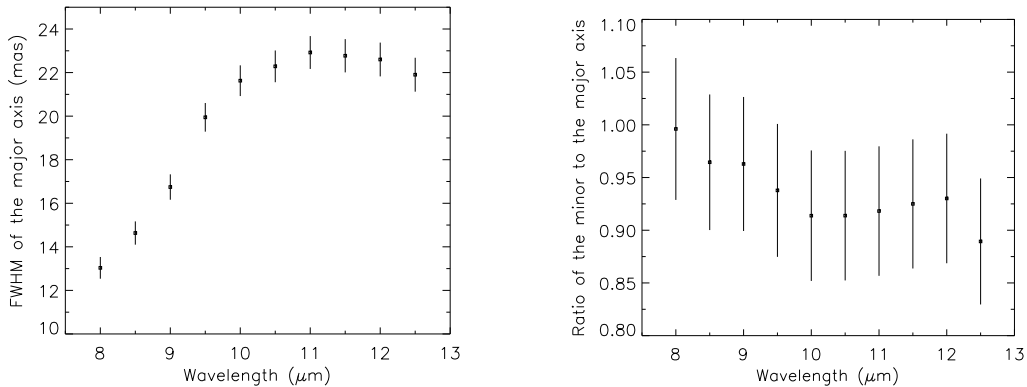


Fig. 6. *Left:* full width at half maximum of the major axis for the best-fit elliptical Gaussian model as a function of wavelength. *Right:* ratio of the minor to the major axis for the best-fit elliptical Gaussian model as a function of wavelength.

ting most of the 8 μm flux, located closer to the photosphere, are the least affected by the White Dwarf, situated far (60 AU; Eyres et al. 2001) from the Mira. At 13 μm , the interferometer probes farther regions that are more affected by the energetic but diluted wind originating in the erupting White Dwarf. We also note that the new baselines, recorded in 2007, range from 60 to 89m whereas the previous ones ranged from 32 to 59m and probed outer regions of the dusty envelope, which were potentially more affected by the White Dwarf. Therefore, we expect the amount of elongation to be higher, as reported in Sacuto et al. (2007), although the poor uv coverage of the 2005-2006 observations does not allow us to estimate it accurately.

These MIDI/VLTI interferometric observations probe a compact region that includes the giant star and its vicinity to a radius of about 30 mas, set by the smallest baselines ($\sim 32\text{m}$). There is a clear requirement for near and mid-IR high dynamics, imaging at the diffraction limit of a 8-10m class telescopes, to probe the inner nebula (0.8'' wide) already studied with the HST (Eyres et al., 2001) and in the radio (Richards et al., 1999).

We are close to the limits of the imaging capabilities of the VLTI used with the two telescopes recombiner MIDI for studying such a complex source. It is also uncertain whether the three

telescopes recombiner AMBER, operating in the near-IR, can provide a substantial increase in information about the morphology of the system, because the (spherical) Mira and hot dust in the inner wind still dominate. A significant improvement in the quality of image reconstruction must await the advent of the second generation instrument MATISSE (Lopez et al., 2006).

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